

DUST DISTRIBUTION IN THE β PICTORIS CIRCUMSTELLAR DISKS

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ABSTRACT

We present 3-D models of dust distribution around β Pictoris that produce the best fits to the Hubble Space Telescope Advanced Camera for Surveys' (HST/ACS) images obtained by Golimowski and co-workers. We allow for the presence of either one or two separate axisymmetric dust disks. The density models are analytical, radial two-power-laws joined smoothly at a cross-over radius with density exponentially decreasing away from the mid-plane of the disks. Two-disk models match the data best, yielding a reduced χ^2 of ~ 1.2 . Our two-disk model reproduces many of the asymmetries reported in the literature and suggests that it is the secondary (tilted) disk which is largely responsible for them. Our model suggests that the secondary disk is not constrained to the inner regions of the system (extending out to at least 250 AU) and that it has a slightly larger total area of dust than the primary, as a result of slower fall-off of density with radius and height. This surprising result raises many questions about the origin and dynamics of such a pair of disks. The disks overlap, but can coexist owing to their low optical depths and therefore long mean collision times. We find that the two disks have dust replenishment times on the order of 10^4 yr at ~ 100 AU, hinting at the presence of planetesimals that are responsible for the production of 2nd generation dust. A plausible conjecture, which needs to be confirmed by physical modeling of the collisional dynamics of bodies in the disks, is that the two observed disks are derived from underlying planetesimal disks; such disks would be anchored by the gravitational influence of planets located at less than 70 AU from β Pic that are themselves in slightly inclined orbits.

Subject headings: formation: disks, planets, dynamics – dust, scattering function

1. INTRODUCTION

Owing to its proximity (19.28 ± 0.19 pc; Crifo et al. 1997), β Pictoris is one of the best studied examples of a Main Sequence (MS) star with a circumstellar dust disk. β Pic was shown to have a substantial mid- and far-infrared excess through observations using the Infrared Astronomical Satellite (IRAS) (Aumann et al. 1984; Aumann 1985; Gillet 1986), believed to be indicative of thermal radiation from ~ 100 K dust orbiting the star. Follow-up imaging by Smith & Terrile (1984) revealed an almost edge-on system with circumstellar nebulosity due to light scattered from the dust around the star. Since then, many investigators have mapped the surface brightness and color of the disk in optical wavelengths (Smith & Terrile 1987; Artymowicz et al. 1989; Lecavelier Des Etangs et al. 1993; Golimowski et al. 1993; Kalas & Jewitt 1995; Heap et al. 2000; Golimowski et al. 2006).

The extensive studies of the disk have also revealed several asymmetries in the scattered light contours produced by the dust surrounding β Pic. These were summarized by Kalas & Jewitt (1995) and include: 1. the size asymmetry - the north-east (NE) extension of the disk stretches further out from β Pictoris than the south-west (SW) extension, 2. the surface brightness asymmetry - the brightness profile along the disk's spine (mid-plane) is broader for the NE than the SW extension, 3. the width asymmetry - the SW extension is thicker than the NE extension, 4. the wing tilt asymmetry - the

two extensions of the disk are not perfectly aligned (i.e. the difference in the position angles of the midplanes of the opposing extensions is less than 180°) and 5. the butterfly asymmetry - isophote curvature is asymmetric across the midplane of the disk, with the asymmetry itself inverted across the minor-projection axis. Several attempts have been made to explain why these asymmetries arise. Kalas & Jewitt (1995) noted that the size and surface brightness asymmetries might be a direct result of the width asymmetry. Since the SW extension of the disk is more vertically extended, the SW midplane might appear fainter than the NE midplane and thus explain why the latter appears to have a larger radial size.

Further insight into the β Pic system has come from optical and mid-infrared observations of its disk. Heap et al. (2000) and Golimowski et al. (2006) imaged the disk with high-resolution HST optical cameras and detected a warp in the inner part of the disk (~ 20 – 100 AU from β Pic) indicating the presence of a secondary disk inclined at $\sim 5^\circ$ to the primary (Heap et al. 2000; Golimowski et al. 2006). Intriguingly, Golimowski et al. (2006) noted that the projected spine of the secondary is aligned with the isophotal inflections (the butterfly asymmetry) reported at large distances from the star. Golimowski et al. (2006) also placed constraints on the brightness profiles along the spines of the primary, secondary and composite (combined primary and secondary) disks. They found that the primary disk has a steeper midplane profile than the secondary at greater distances from the star and that the composite profile is best described with four power laws, in contrast to previous modelling efforts that employed two power laws (eg. Kalas & Jewitt 1995). To probe the regions of β Pic's

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disk inaccessible with optical instruments, Wahhaj et al. (2003) imaged the disk in the mid-infrared ($17.9\ \mu\text{m}$). They reported a set of rings embedded in β Pic's disk whose radii vary from 14 to 82 AU and whose configuration suggests the presence of multiple planets. However, the ring structure is uncertain, since Golimowski et al. (2006) failed to observe them in scattered light.

Strong motivation for the study of dust distribution in disks is spurred by the fact that such a study reveals important clues about the presence of unseen larger bodies (planetesimals and planets). For example, it was suggested that the HST warp (Heap et al. 2000) is caused by the presence of a giant planet on a slightly inclined orbit less than 50 AU from the primary (ex. Mouillet et al. 1997). This prediction was possibly vindicated recently by the announcement of an $8M_{\text{Jup}}$ candidate companion separated by 8 AU from β Pic, which, if confirmed, would make it one of the first planets imaged orbiting an A type MS star, and the planet imaged with the smallest projected physical separation to-date (Lagrange et al. 2009).

The system's many peculiarities have yet to be adequately explained, and thus continued studies of the dust distribution are a necessity that might lead to the discovery of additional companions. Given the possibility of dust migration (Takeuchi & Artymowicz 2001) and/or dust avalanches (Grigorieva et al. 2007), the dust and the planetesimal/planet distributions may not be the same. However, detailed modeling of micron-sized dust including its dynamical interaction with radiation and dust-dust collisions, is a prerequisite for understanding the distribution of larger bodies. Our work provides the first step for such prospective modeling, the description of the dust distribution.

In this paper we present the best-fit two-disk model to the images obtained by Golimowski et al. (2006). In section 2 we discuss the model assumptions as well as the ACS images that are used for the fitting procedure. We dedicate section 3 to the results and section 4 to the discussion and analysis.

2. MODELING AND DATA

2.1. Multi-Parametric Model

We use an axisymmetric multi-parameter model to describe the distribution of the dust surrounding β Pic. Its simplicity and the fact that it has been successfully applied in the past on more than one occasion (eg. Artymowicz et al. 1989; Kalas & Jewitt 1995) allow one to quickly proceed from the numerical development stage to testing the model on astronomical data. The improvements on the previous applications of the multi-parametric model include the introduction of an additional axis of rotation (which allows us to model disks whose axis can point in any direction), implementation of a χ^2 minimization algorithm (in this case a Markov chain Monte Carlo method) and the addition of the secondary disk which we fit simultaneously with the primary. Furthermore, recent Golimowski et al. (2006) observations have produced the most photometrically accurate maps to date which allow one to probe dust distribution in all directions (specifically at large scale heights), allowing one to introduce additional constraints on the model.

In order to describe the dust distribution, we follow Artymowicz et al. (1989). We assume that both disks are axisymmetric. In cylindrical co-ordinates (r, z, θ) then,

the scattering cross section per unit volume is given by:

$$\frac{d\tau}{ds} = \frac{\tau(r)}{W(r)} e^{-(|z|/W(r))^p} \quad (1)$$

This form arises as a result of the vertical exponential drop off of the dust density. The rate of this drop off is determined by the parameter p . $W(r)$ is the width profile:

$$W(r) = A_0 R_m \left(\frac{r}{R_m} \right)^\gamma \quad (2)$$

This profile guarantees that the disk is thicker at greater radii than at smaller radii. If γ , the flare index, is more than 1, the disk will flare. The scale height A_0 ensures that the disk has a specific thickness to radius ratio at $r = R_m$, the location of the power law break describing the vertical optical thickness.

The vertical optical thickness is proportional to the function $\tau(r)$. $\tau(r)$ has the following dependance on distance from the centre of the star:

$$\tau(r) = \frac{\eta}{\sqrt{(r/R_m)^{-2\alpha} + (r/R_m)^{-2\beta}}} \quad (3)$$

The inner radial index α and the outer radial index β suppress the vertical optical thickness below and above the characteristic radius R_m . Only two indices are used since the midplane surface brightness profile of the primary disk is known to have a single break and thus requires two power laws to describe it (Artymowicz et al. 1989; Kalas & Jewitt 1995; Golimowski et al. 2006).

To compare the brightness contours produced by our model with images of β Pic's disk, we convert the cylindrical symmetric dust distribution into the observed brightness isophotes in the following manner. Light intensity along the line of sight through the disk is given by:

$$I(x, y) = \int \frac{\frac{d\tau}{ds}(x, y, l) f(\theta) dl}{4\pi(r^2 + z^2)} \quad (4)$$

Here x and y are the pixel co-ordinates of an image and l is the depth along the line of sight. In the above equation, flux entering a unit volume having co-ordinates (r, z) is $L_\nu/[4\pi(r^2 + z^2)]$, where L_ν denotes the luminosity of β Pic at frequency ν . This luminosity, along with the distance to the β Pic, is hidden inside our normalization parameter η . $f(\theta)$ represents the phase function and is described in section 2.2. The formulae for r , z and θ as functions of x , y and l follow from the appropriate co-ordinate transformations.

$$r = [(l \cos i + y \sin i)^2 + x^2]^{1/2} \quad (5)$$

$$z = y \cos i - l \sin i \quad (6)$$

$$\theta = \cos^{-1} \left(-l/(x^2 + y^2 + l^2)^{1/2} \right) \quad (7)$$

The theoretical image obtained via the described transformation will show a disk that is aligned with the major projection axis. Angle ϕ is introduced so that this image can be rotated around the line-of-sight until it is aligned with a disk, which may not lie along the major projection axis (which is certainly the case for the secondary).

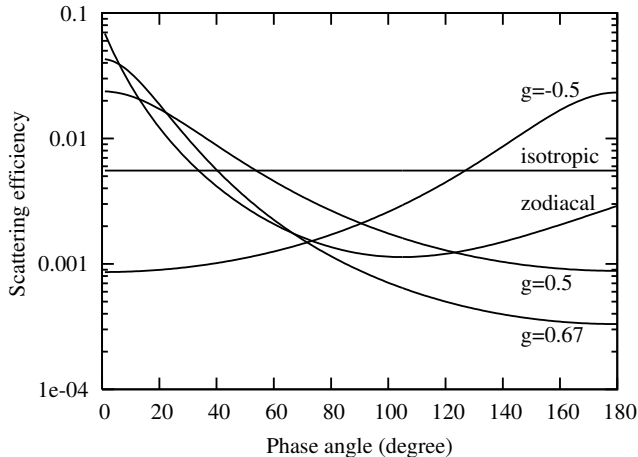


FIG. 1.— The scattering efficiency as a function of phase angle is plotted for empirical zodiacal (Leinert et al. 1976) and several Henyey-Greenstein fits (Henyey & Greenstein 1941), including isotropic scattering ($g=0$).

2.2. Phase Functions

In order to obtain a theoretical image, one must also include a phase function [$f(\theta)$ in equation 4] which describes how the light is scattered in different directions compared to the line-of-sight by the orbiting dust¹. For our simulations we used two scattering profiles, the empirical zodiacal (Leinert et al. 1976) and the Henyey-Greenstein (Henyey & Greenstein 1941). The former profile describes light scattered by the midsized particles (10-100 μm) with rough surfaces (Schiffer & Thielheim 1982) and it strongly favors forward scattering. The empirical zodiacal is proportional to:

$$f(\theta) = 0.3(0.2 + \theta/2)^{-3} + 1.4(\theta/3.3)^4 + 0.2. \quad (8)$$

The Henyey-Greenstein scattering profile is an analytical phase function commonly used to describe light scattered by small grains:

$$f(\theta) = \frac{1 - g^2}{4\pi(1 + g^2 - 2g \cos \theta)^{3/2}}, \quad (9)$$

where $-1 < g < 1$ is the profile asymmetry parameter. Positive values of g produce forward scattering, negative backward while $g = 0$ gives us the isotropic scattering. In Fig.1 we present the zodiacal light and several Henyey-Greenstein phase functions. In the simulations (as well as Fig.1) the phase functions are normalized so that their integrals over the full solid angle are equal. This is done so that a meaningful comparison can be made between η parameters obtained from different simulations.

2.3. Data

In order to constrain the two-disk model of β Pic's dust distribution, we use the Golimowski et al. (2006) ACS Hubble Space Telescope (HST) images. The ACS' excellent optics and HST's well studied PSF allow one to produce images whose quality (photometric accuracy and morphological details) are unparalleled by any other hardware combination currently in operation. Furthermore, the ACS' wide field-of-view (FOV) [$\sim 29'' \times 26''$],

its ability to access large scale heights [$\sim 1.5''$ circular coronagraphic mask only blocks out the innermost 30 AU] and its ability to use three filters [F435W(B), F606W(broad V) and F814(broad I)] to probe chromatic dependence introduce additional constraints on our model that would have been impossible with previous observations.

To discriminate between the features in the image associated with the disk and those intrinsic to the coronagraphic PSF of the disk and β Pic itself, Golimowski et al. (2006) observed the disk at two roll angles and obtained images of a star with similar colors (α Pic) prior to every exposure. Following the appropriate subtractions, the combined images were deconvolved with synthetic PSFs produced by the Tiny Tim software package². Throughout the reduction process Golimowski et al. (2006) keep track of errors (both random and systematic) associated with each pixel. This allows them to produce meaningful error maps that include both the systematic and statistical uncertainties in their data; we use these error maps for our χ^2 minimization. The number of data points in these images once the coronagraph (inner 30 AU) is masked out is 563136.

2.4. Markov chain Monte Carlo

We employ Markov chain Monte Carlo (MCMC; Ford 2005) fitting as described for our purposes in Croll 2006. MCMC fitting is a computationally efficient method to perform a full Bayesian analysis of complicated problems where one is required to fit a great number of often correlated parameters. MCMC fitting effectively samples the posterior parameter distribution by performing an n-step intelligent random walk around the K -dimensional parameter space of interest, while recording the χ^2 at each point in the intelligent random walk. In addition, MCMC fitting allows one to explore correlations in one's fitting-parameters - the end result is that one is able to return realistic best-fit and error estimates even if the parameters of interest are correlated. MCMC fitting is thus well-suited to fit the Golimowski et al. (2006) observations of the β Pic disk with our model as it features a large number of parameters ($K=18-20$ parameters for the two-disk case) a select number of which display modest correlations.

Flat priors in all parameters are used. We run our MCMC chains until the Gelman & Rubin (1992) statistic is close to unity for all parameters. To determine our best-fit and $1-\sigma$ uncertainties, we use the marginalized likelihood method as described in Croll (2006).

2.4.1. Accounting for the large number of data points

Given the large number of data-points (>500000) in the ACS images of β Pic and the impressive accuracy of this data minute changes in one's model will result in a statistically significant increase in χ^2 . These small increases in χ^2 indicate a statistically significant poorer fit in the case that one's model is an accurate representation of the physical reality of the disk in this case. However, in our case it should be noted that even in the best-case scenario our model is likely only an approximation of the actual physical makeup of the β Pic disk; even

¹ We assume the same optical properties of dust everywhere.

² For a more detailed discussion of the reduction procedure, we refer the reader to Golimowski et al. (2006)

our best-case scenario cannot account for the numerous asymmetries in the disk (e.g. Figure 5). Furthermore, although the error estimates provided by Golimowski et al. (2006) attempt to include both the systematic and statistical uncertainty in the ACS data, they could fail to account for correlations in the data that could add an extra systematic component. For these reasons we feel that a simplistic χ^2 could underestimate the true uncertainties in our fitted parameters. For these reasons we have scaled up the uncertainties in our fitted parameters by a factor of 10, as we believe these give a more accurate indication of the true uncertainty in our data.

3. RESULTS

In the following sections we discuss optimal parameters for a single and two-disk dust distributions as well as several scattering phase functions. Since both the dust distribution and the phase function are responsible for the scattered light profile, both components are investigated here and modelled with our simulations. However, in order to minimize the number of free parameters, we first explore the parameter space using a zodiacal light phase function for scattering. In section 3.3 we expand on our work by exploring other phase distributions.

3.1. One Disk vs. Two Disk Model

To test both the fitting procedure and the goodness-of-fit of our two-disk fits, we first constrained the parameters describing the dust distribution via a single disk and a zodiacal light scattering profile. The zodiacal light profile has been used in the past (eg. Kalas & Jewitt 1995) and was chosen in order to minimize the number of free parameters. Table 1 lists these results as well as those obtained in the past using similar axisymmetric models. We do not note any significant disagreements and suggest that the minor differences between our and past results might be due to several factors including: 1. that previous fitting efforts were performed with simpler models (eg. Kalas & Jewitt 1995), 2. that the previous data used for modelling efforts was of poorer quality, 3. that various coronagraphic implementations employed in previous observations did not allow the observer to reliably probe large scale heights, 4. the usage of different phase functions and lastly 5. that past modellers applied trial and error methods to explore the parameter space instead of more sophisticated techniques.

After we obtained the best single disk (9 parameter) fit, we ran our fitting procedure using two disks (18 parameters). The best-fit results for the one and two-disk models are listed in Table 2 along with the associated errors on each parameter - the MCMC probability distributions of the 18 parameter 606 fit are presented in Figure 3. The minimum reduced χ^2 for the single and two-disk fits are 1.83 and 1.18 respectively. In addition to the improvement in the reduced χ^2 , we present other qualitative improvements. In Figure 4 we display the isophotal maps of our best-fit models, which shows that the two-disk model's isophotes better trace the observed isophotes. The isophotal maps also reveal that two disks with axisymmetric dust distribution and non-isotropic scattering profiles can produce asymmetric brightness contours if they are inclined to each other and the line-of-sight (LOS) at non-zero angles. Degree of this asymmetry is further discussed in section 4.1.

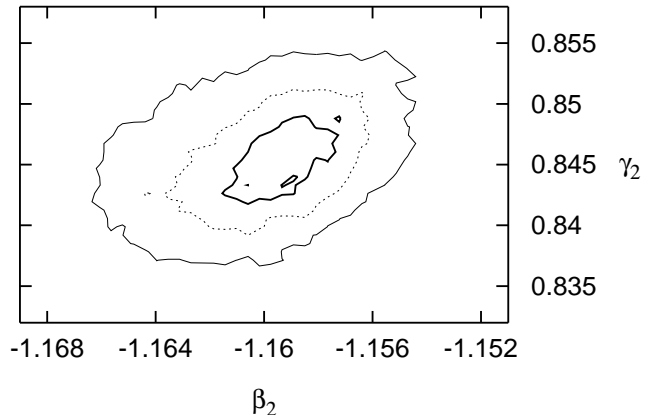


FIG. 2.— A plot of the correlation between the radial index (β_2) and the flare index (γ_2) displaying the 33% (solid thick curve), 68% (dotted curve) and the 95% (solid thin curve) credible regions for the 18-parameter 606 ACS data. While some correlation between these parameters is expected as they both describe light scattered at larger radii, the two are still well constrained.

The total residuals following the removal of the one and then two-disk models from the data are presented in Figure 5. Aside from producing overall smaller residuals, the two-disk model residuals are also symmetric, revealing the main warp at (65-85) AU as well as an additional structure at (110-150) AU which is similar to the rings observed by Wahhaj et al. (2003). We further discuss these structures in section 4.1. In Figure 6 we present brightness profile cuts along both the NE and the SW extensions of the composite disk, demonstrating once again the improvements gained from employing the two-disk model. For these reasons we believe the β Pic disk is much more accurately described by our two axisymmetric disk model¹.

We also investigate for possible correlations between the various parameters of our model. The most significant correlation is between the secondary's outer radial index (β_2) and the flare index (γ_2). This correlation is shown in Figure 2 for the 18-parameter F606W data set. Other parameters were not found to be notably correlated.

3.2. Chromatic dependence of modeled parameters

Prior to the Golimowski et al. (2006) observations of β Pic, it was widely accepted that the orbiting dust was composed of neutrally scattering grains with sizes larger than several microns (ex. Lecavelier Des Etangs et al. 1993). With the higher photometric precision of the ACS, Golimowski et al. (2006) analyzed the disk's colors with three filters [a narrow B-filter (435 nm), a broadband V (606 nm), and a broadband I (814 nm)] and concluded that this assumption was wrong. Complicating matters further, their observations showed asymmetries in the flux ratios taken at different wavebands about the projected major and minor axis of the disk, which imposes additional constraints on the models and suggests that a

¹ Bayesian model comparison over the one-disk model while taking into accounts the extra degrees of freedom of the two-disk model. It was not applied here due to the clear quantitative and qualitative improvements of the two-disk model over the one-disk model

TABLE 1
SINGLE DISK PARAMETER COMPARISON

Parameter	Kalas & Jewitt (1995) ^{a,c}	Artymowicz (1998) ^{b,c}	This paper
Inclination	$2^\circ < i < 5^\circ$	$i = 1.3^\circ$	$i = 2.3^\circ$
Vertical distribution	$0.7 < p < 2.0$	$p = 0.7$	$p = 0.5$
Radius of the power law break	$R_m = 100$ AU	$R_m = 120$ AU	$R_m = 102$ AU
Inner radial index	N/A	$\alpha = 2.0$	$\alpha = 2.0$
Outer radial index	$-3.4 < \beta < -2.8$	$\beta = -3.0$	$\beta = -1.8$
Flare index	$1.1 < \gamma < 1.6$	$\gamma = 0.75$	$\gamma = 1.5$
Scale height	$0.05 < A_0 < 0.10$	$A_0 = 0.055$	$A_0 = 0.02$

^a Parameter values converted from Kalas & Jewitt (1995) to match our model values.^b Personal communiqué. Results obtained by fitting Heap et al. (2000) HST STIS images.^c Fits obtained by trial and error.TABLE 2
9, 18 AND 20 PARAMETER FITS

Parameter	9 parameter 606 χ^2 minimum	9 parameter 606 MCMC best-fit	18 parameter 606 χ^2 minimum	18 parameter 606 MCMC best-fit	20 parameter 606 χ^2 minimum	20 parameter 606 MCMC best-fit
χ^2	1.0390×10^6	$1.0394^{+0.0002}_{-0.0028} \times 10^6$	6.664×10^5	$6.664^{+0.006}_{-0.002} \times 10^5$	6.6×10^5	$6.6^{+0.1}_{-0.2} \times 10^5$
Reduced χ^2	1.8451	$1.8458^{+0.0004}_{-0.0049}$	1.1833	$1.1834^{+0.0010}_{-0.0003}$	1.17	$1.17^{+0.02}_{-0.04}$
i_1 ($^\circ$)	2.261	$2.253^{+0.194}_{-0.009}$	0.11	$0.12^{+0.06}_{-0.06}$	0.12	$0.12^{+0.03}_{-0.04}$
p_1	0.455	$0.458^{+0.006}_{-0.014}$	0.84	$0.85^{+0.01}_{-0.02}$	0.746	$0.752^{+0.061}_{-0.002}$
ϕ_1 ($^\circ$)	1.015	$1.035^{+0.010}_{-0.183}$	0.506	$0.505^{+0.028}_{-0.006}$	0.48	$0.50^{+0.07}_{-0.18}$
α_1	2.04	$2.05^{+0.11}_{-0.08}$	2.73	$2.73^{+0.12}_{-0.08}$	4.8	$4.7^{+0.5}_{-0.2}$
β_1	-1.84	$-1.84^{+0.03}_{-0.02}$	-2.89	$-2.89^{+0.07}_{-0.04}$	-2.8	$-2.8^{+0.2}_{-0.2}$
γ_1	1.51	$1.51^{+0.06}_{-0.02}$	1.05	$1.05^{+0.04}_{-0.03}$	1.16	$1.18^{+0.04}_{-0.20}$
i_2 ($^\circ$)	n/a	n/a	6.312	$6.312^{+0.033}_{-0.010}$	2.6	$2.7^{+0.3}_{-0.8}$
p_2	n/a	n/a	0.464	$0.465^{+0.024}_{-0.008}$	1.1	$1.0^{+0.7}_{-0.2}$
ϕ_2 ($^\circ$)	n/a	n/a	3.871	$3.871^{+0.003}_{-0.003}$	5.3	$5.4^{+0.3}_{-0.4}$
α_2	n/a	n/a	1.6	$1.6^{+0.3}_{-0.3}$	9	9^{+1}_{-3}
β_2	n/a	n/a	-1.16	$-1.16^{+0.02}_{-0.03}$	-0.25	$-0.26^{+0.11}_{-0.04}$
γ_2	n/a	n/a	0.84	$0.84^{+0.05}_{-0.02}$	1.59	$1.57^{+0.07}_{-0.06}$
R_{m1} (AU)	102	102^{+1}_{-2}	112.7	$112.7^{+0.9}_{-1.2}$	112	112^{+2}_{-2}
R_{m2} (AU)	n/a	n/a	71	71^{+3}_{-3}	63.5	$63.1^{+11.1}_{-0.7}$
η_1	3.96×10^{-10}	$4.02^{+0.04}_{-0.28} \times 10^{-10}$	5.60×10^{-10}	$5.62^{+0.06}_{-0.06} \times 10^{-10}$	5.63×10^{-10}	$5.66^{+0.08}_{-0.25} \times 10^{-10}$
η_2	n/a	n/a	2.0×10^{-10}	$2.0^{+0.2}_{-0.1} \times 10^{-10}$	2.41×10^{-10}	$2.40^{+0.01}_{-0.17} \times 10^{-10}$
A_{01}	0.018	$0.018^{+0.002}_{-0.002}$	0.0489	$0.0492^{+0.0010}_{-0.0012}$	0.0429	$0.0435^{+0.0043}_{-0.0009}$
A_{02}	n/a	n/a	0.051	$0.052^{+0.012}_{-0.005}$	0.056	$0.057^{+0.004}_{-0.008}$
g_1	n/a	n/a	n/a	n/a	0.637	$0.635^{+0.006}_{-0.004}$
g_2	n/a	n/a	n/a	n/a	0.849	$0.848^{+0.005}_{-0.007}$

^a As discussed in §2.4.1 the formal errors have been scaled up by a factor of 10.

true description of the dust distribution will be incomplete until wavelength-dependent models are developed.

Unfortunately, our multi-parametric dust distribution model is not wavelength-dependent and neither are the phase functions we use³. In addition, any modifications would introduce additional parameters and thus slow down the search through the parameter space. Considering these points and the already low reduced χ^2 we have been able to obtain from our wavelength-independent model, we chose not to increase the complexity of the model.

There are some benefits in our fitting the same model to data observed at different wavelengths. Fitting to images of the disk at different wavelengths and comparing

³ zodiacal light function is derived from particles scattering sunlight which are much larger than the wavelengths they scatter and thus one can think of it as color-neutral scattering function

the derived parameters can reveal whether the scattering in the β Pic disk is wavelength dependent and even put some constraints on the optical properties of dust. Conversely, assuming the scattering in the β Pic disk to be color-neutral and comparing the results from fits using different wavelengths, we can obtain an independent estimate of the robustness of our fit.

Table 3 lists the best fits obtained for the three wavebands. Overall there is a good agreement between the parameters, especially for the values describing the primary. Several of the parameters describing the secondary [the inner radial index α_2 , radius of the power law break R_{m2} and the scale height index A_0] show a much wider range than those corresponding to the primary disk. For instance the apparent increase of α_2 in the shorter wavelength bands might suggest that there is substantial reddening in the inner part of the secondary. However, we

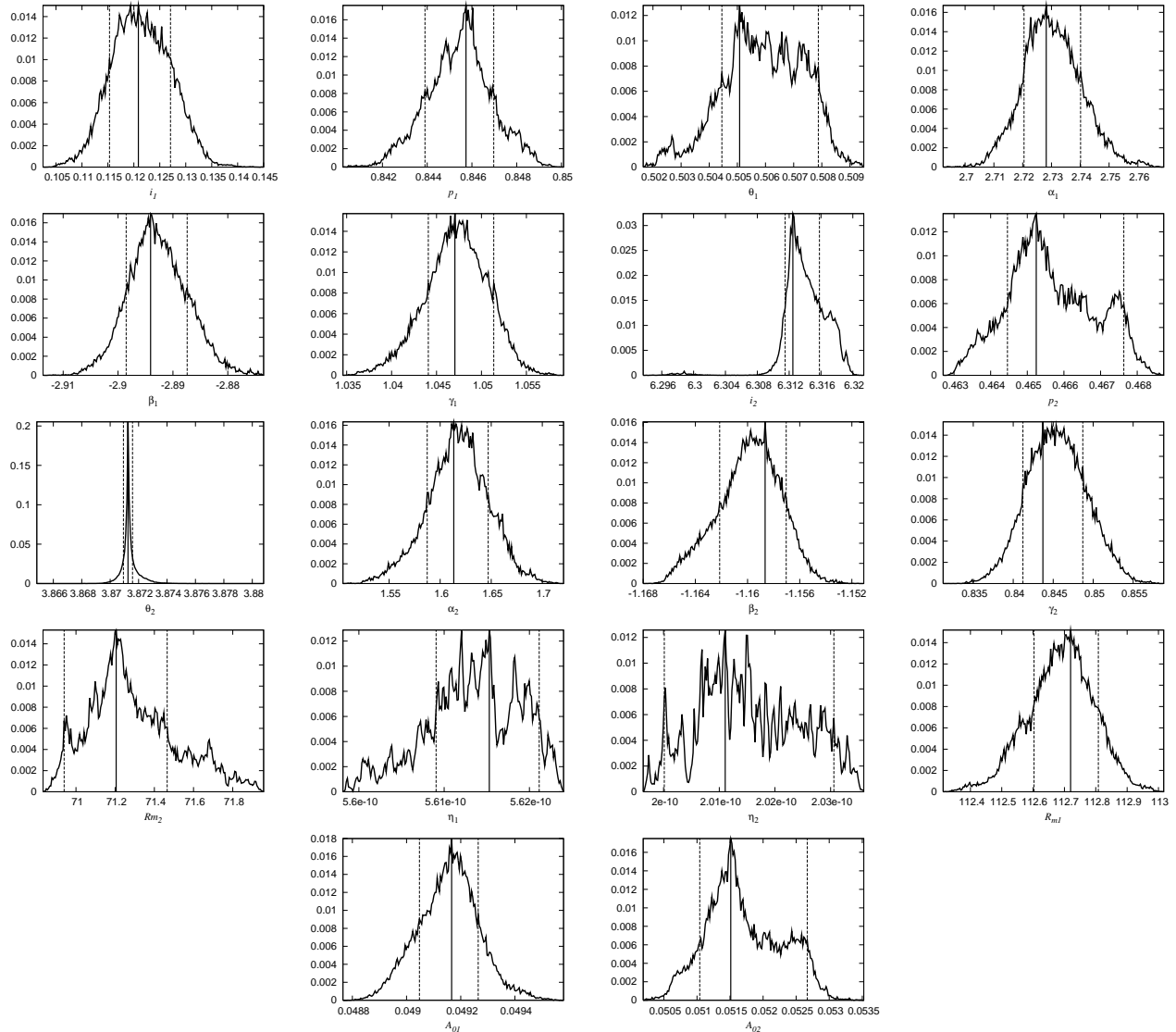


FIG. 3.— The marginalized likelihood probability distributions of our 18-parameter 606 MCMC chains for each of the fitted parameters. The best-fit value (the peak of the distribution) and the 68% credible regions are given by the solid and dashed vertical lines.

also note that any errors in the subtraction near the coronagraphic mask would effect α_2 the most and thus we believe that this result should be viewed with some skepticism.

We also use the models produced for the three filters to constrain the inclination of the secondary to the primary. The fits for the locations of the spines of both disks (as well as the composite) are shown in Figure 7. Table 4 lists the corresponding inclinations of the NE and SW extensions measured counterclockwise to the major projection axis. The value for the angular separation is of particular interest since this value has been previously reported. Our result, $3.2 \pm 1.3^\circ$, is consistent with the $\sim 5^\circ$ value reported by Golimowski et al. (2006).

3.3. Phase Function

A topic of importance that has largely been ignored until now is the appropriateness of different phase functions to model the scattering of light in the β Pictoris disk. Since an observed brightness contour is a function of both the distribution of scatterers as well as the phase

function that describes how these particles scatter light, a discussion of the latter is necessary. The zodiacal scattering profile (Fig. 1) was initially chosen for two reasons: 1. previous successful modeling with this function (eg. Kalas & Jewitt 1995) and 2. its empirical formula does not require any additional parameters. As it was already noted, using this profile has allowed us to obtain a very low reduced $\chi^2 = 1.18$. However, the question remains: how valid is this choice? One way to answer this question is to employ other phase functions. In section 2.2 we introduced the Henyey-Greenstein function along with the zodiacal scattering profile. The former is particularly interesting because it can be modified through the scattering asymmetry parameter g to produce many profiles, several of which are shown in Fig. 1. To obtain the best fit, we add the scattering parameter to the 18 parameters describing dust distribution with two disks and run the MCMC algorithm until the different chains converge. Intriguingly, the 20 parameter fit does not offer noticeable improvements in the reduced χ^2 (~ 1.2 for both 18 and 20 parameters). What's more, the values

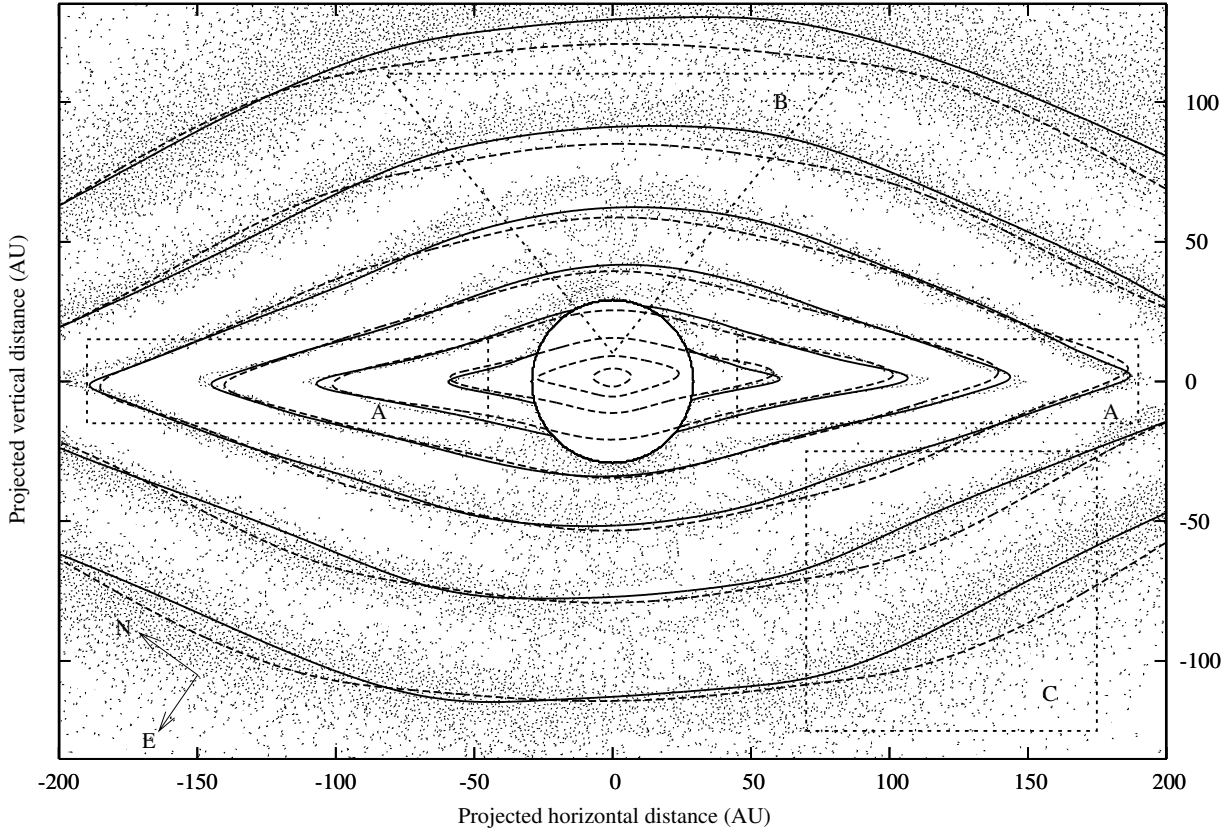


FIG. 4.— Isophotal map of the light scattered by the dust orbiting β Pic (dotted regions). Also shown are the single (dashed line) and the two-disk (solid line) model scattered light profiles. The brightness isophotes are quoted relative to β Pic, with the isophotes decreasing from $1 \times 10^{-6} L_{\text{disk}}/L_{\beta\text{Pic}}$ with interval spacing equal to 0.5 in log base 10 units. From the overall low reduced $\chi^2 = 1.84$ of a single disk fit, one expects a reasonable fit for the observations. This is evident in the isophotal maps. The improvements offered by the two-disk model are shown in several regions: along the spine (A), forward scattering areas originating from larger scale heights (B) and the overall improvement in tracing the scattered light isophotes (C).

of the scattering asymmetry profile favored by the parameter search ($g_1 = 0.64$ and $g_2 = 0.85$ for the two disks) produce a forward scattering profile that largely resembles the zodiacal scattering function; furthermore the uncertainty in $g_1 = 0.64$ and $g_2 = 0.85$ are small (see Table 2). The disagreement lies in the slightly higher amount of back scatter produced by the latter, although this is not significant since most of the light comes from forward scattering. In Table 2 we present the results of our analysis of the two-disk run that allowed for varying scattering asymmetry parameters.

4. DISCUSSION

4.1. Morphology of the Composite Disk and its Components

The overall shape of the brightness contours produced by the two-disk model follows that of the light scattered by β Pic's dust. This is evident from both the isophotal maps (Figure 4) and low $\chi^2 \sim 1.2$ fits for the F606W filter images. Similar maps and reduced χ^2 are recorded in other filters as well. The brightness contours of the two components (the primary and the secondary) is presented in Figure 8 and the relative contribution of each component to the total scattered light is shown in Figure

9.

The most striking feature of these fits is that the secondary contributes more light at larger radii and scale heights than the primary. The immediate implication is that the dust content of the secondary disk is not localized to the inner regions of the system and that it extends at least out to ~ 250 AU. We noted earlier that Golimowski et al. (2006) found a steeper power law fit along the spine of the primary beyond ~ 120 AU than the secondary. Their result suggests that the secondary will extend further out, at least radially, unless the fit is not appropriate beyond ~ 250 AU and is in line with our conclusions. It is also interesting that to describe the brightness profile along the midplane of the composite, Golimowski et al. (2006) use 4 power laws with power law breaks at 69, 117 and 193 AU. Our two-disk model offers an explanation for the locations of the breaks. The 69 and 117 AU breaks correspond closely to the midplane profile breaks of the secondary (71 AU) and the primary (113 AU). The last power law break according to our model is a result of primary's dust density falling to levels equal to the secondary's density along the line of sight and the contour transitioning to follow the shape of the secondary.

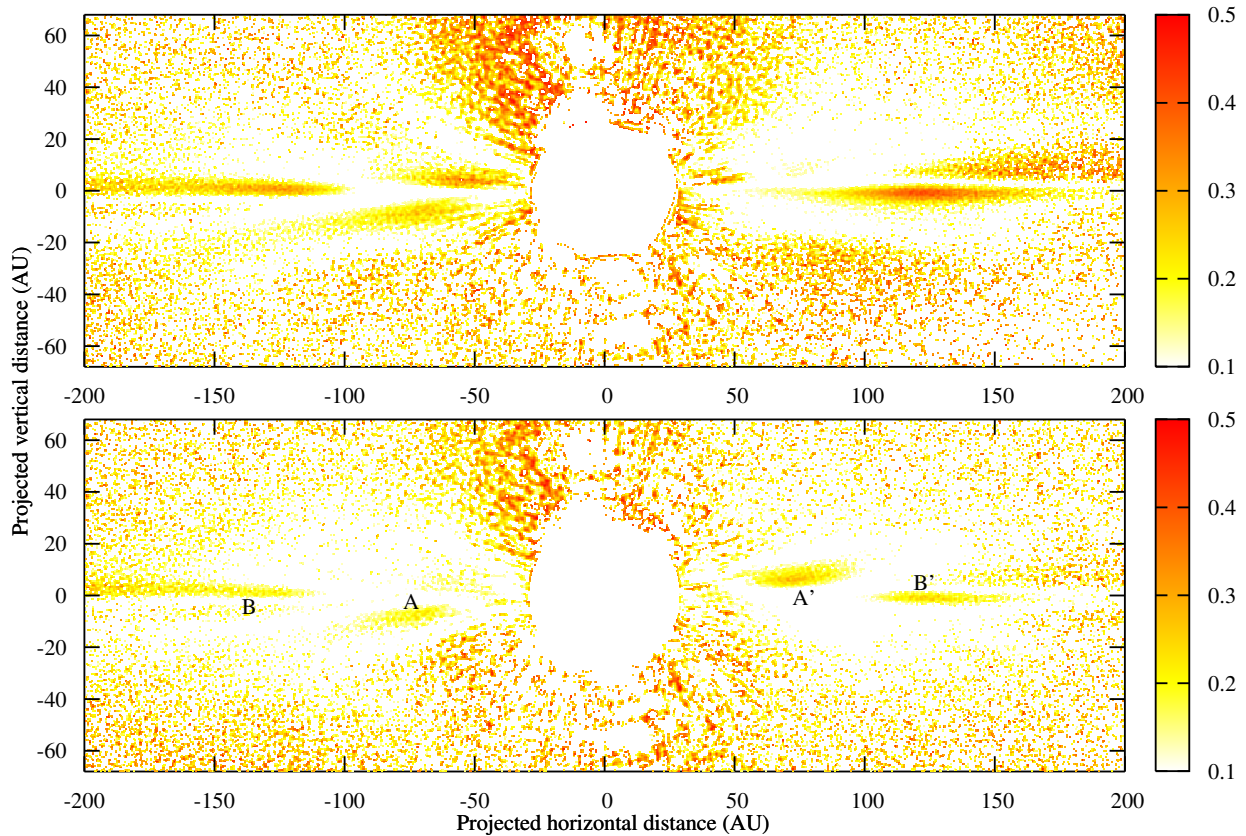


FIG. 5.— The residual of the 606 ACS data following the removal of the scattered light produced by the single (top panel) and the 2 disk (bottom panel) models, presented relative to that of the composite disk. The second residual is not only cleaner, but is also symmetric following a 180° rotation. The identified ring-like structures (A and A', B and B') are similar to those identified by Wahhaj et al. (2003), albeit we are not certain whether their origin is real or failure of the model to properly describe the dust distribution.

In order to test for possible disagreements with previous ground base observations of the disk which extend out to 800AU (e.g. Kalas & Jewitt 1995) and do not show two separate components, we extend our model to the same radius. Due to the large scale height of the secondary disk in our model, we find that the composite still appears as a single disk at large radial separations (out to ~ 800 AU). Furthermore, the primary makes a significant contribution to the total light profile along the spine, having the effect of aligning the composite's spine with its own. Thus, the composite does not show an inclination to the primary at 800 AU.

The impact of such an extended secondary on dust dynamics in the system is discussed in section 4.3.

Figure 4 also reveals that the brightness isophotes of the two-disk (composite) model are spaced more widely along the northwestern semiminor axis, as can be seen from the greater extension of the isophotes in the northwest direction. Golimowski et al. (2006) attribute this spacing to the northwest side of the disk being tipped nearer to earth. All of our two-disk models agree that the spacing is a result of the secondary's inclination to the line of sight ($i_2 = 6.0 \pm 1.0^\circ$ vs $i_1 = 0.0 \pm 0.1^\circ$ for the primary which is almost entirely edge on). This is highlighted in Figure 8 which maps the isophotes of the scattered light of the primary and the secondary showing

that the secondary's isophotes are more widely spaced in the NW than the SE, whereas the primary fails to show this asymmetry.

It is also interesting to compare brightness contours of the primary and the secondary at different wavelengths. In Figure 10 we present color ratios of the F814W and F435W filters for the primary and the secondary. Golimowski et al. (2006) noted that the composite disk ratio shows asymmetry along both the major and the minor projection axis. The latter of these was attributed to dust grains scattering blue light more isotropically than red light. Our model re-creates this asymmetry (Figure 10-top panel). The ratios of the primary at different filters (Figure 10-middle panel) do not show this asymmetry, which is expected from an axisymmetric edge-on disk. Our models suggest that it is the secondary which is responsible for the mentioned minor-axis asymmetry (which is expected considering its much larger inclination to the line-of-sight as well as the primary, which was used to define the major and the minor axis).

Lastly, we comment on the residuals obtained from the subtraction of the composite from the original image (Figure 4). As has already been noted, the residual is symmetric resembling rings observed by Wahhaj et al. (2003) in the mid-IR. The radii of these remnants are

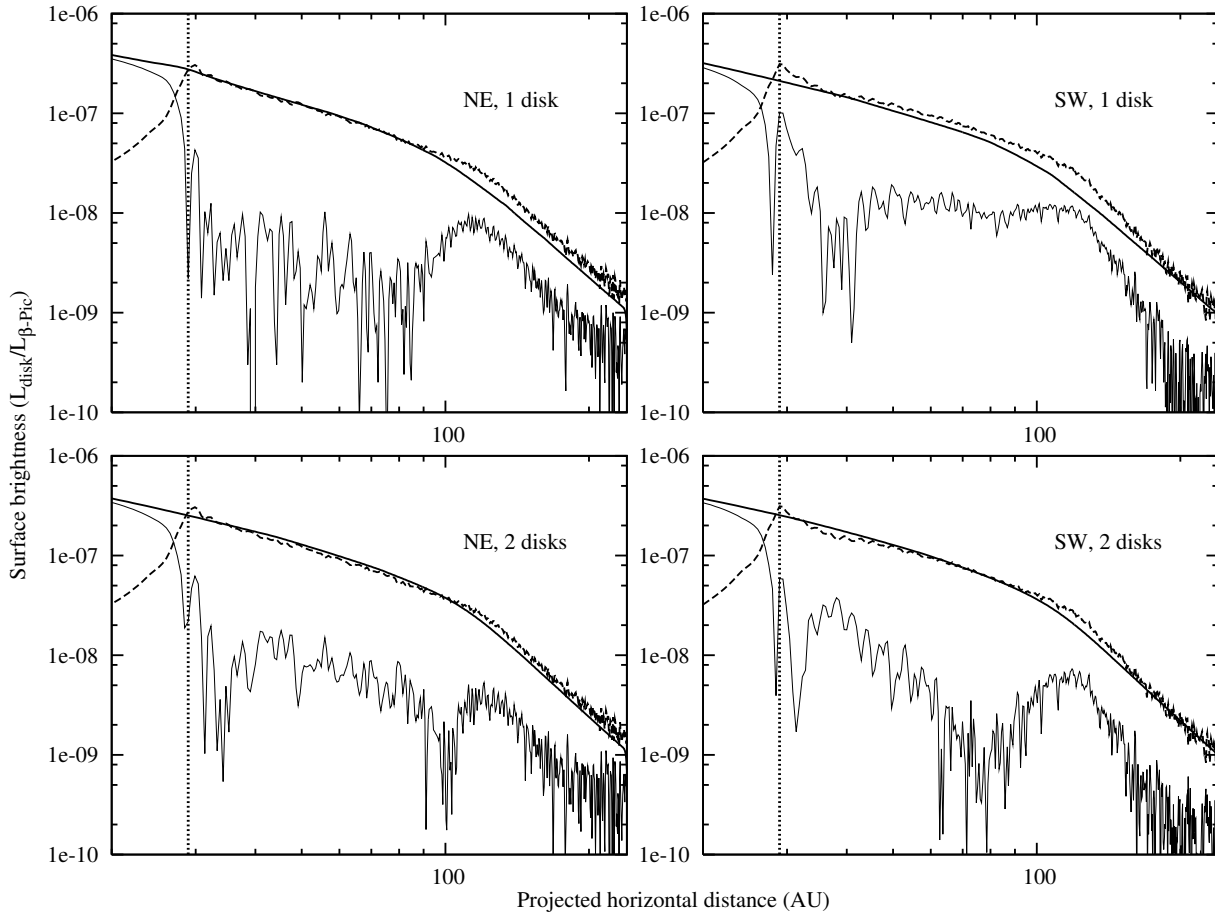


FIG. 6.— Scattered light profile along the semimajor projection axis (the x-axis of Fig. 4). Top panels show the northeast and the southwest extensions (dashed lined) along with profiles produced by single disk model (thick solid line) and the residual (thin solid line) of these two. Bottom panels show the same for the two-disk model. Improvements in the fit can be seen by comparing the two-disk fit with the 1 disk fit. The dotted vertical line at 30 AU represents the extent of the coronagraphic mask.

(65 – 85) AU and (110 – 150) AU, with the inner ring-like structure having a similar radius and orientation as the most outer ring detected by Wahhaj et al. (2003) ($R=82 \pm 2$ AU, $i=-2 \pm 2^\circ$) and coinciding with the HST warp detected by Heap et al. (2000). Since this structure is detected in the visual and the mid-IR regimes, we do not doubt its existence. Furthermore, if the inner “ring” is indeed a structure separate from the two disks, it is not surprising that our simple axisymmetric descriptions fail to account for its presence. Unlike the inner “ring”, the outer structure has not been previously observed. This raises a possibility that it is a byproduct of the fitting procedure. For example, the two-disk fit might be overcompensating for the inner warp by slightly inclining the secondary disk and increasing the overlapping region. Our models fail to identify the 52 AU ring, confirming the absence reported by Golimowski et al. (2006).

4.2. Asymmetries

High resolution images of dust surrounding β Pictoris obtained by Golimowski et al. (2006), with their excellent morphological detail and photometric accuracy, impose significant constraints on models. In addition, ratios of ACS images before and after deconvolution show that dust around β -pic is distributed in two separate disks inclined at $\sim 5^\circ$ to each other. A good model should there-

fore be able to reproduce the morphology with two nearly edge-on disks included by a small angle. As shown already, the low reduced χ^2 and the brightness isophotes shown in Figure 4 demonstrate the success of the MCMC algorithm we have employed to sweep through the applicable parameter space. The small inclination of the primary to the secondary ($3.2 \pm 1.3^\circ$) and the line of sight ($i_1 = 0.0 \pm 0.1^\circ$) adds additional credibility to our fits. However, no model of β Pic’s dust can be said to be successful unless it recreates at least several asymmetries mentioned in the literature, such as the wing tilt and the butterfly asymmetry. Below we discuss the known asymmetries and the degree to which our model recreates them.

Butterfly Asymmetry: The asymmetry first identified by Kalas & Jewitt (1995) refers to the asymmetric curvature of the isophotes across the spine of the composite disk and inversion of this asymmetry across the minor axis. Figure 4 confirms that the two-disk model does a good job of re-creating this feature. Also, since the primary disk is edge-on and nearly identically aligned with the major-projection axis, it is the extended secondary that is responsible for this feature.

Wing-Tilt Asymmetry: Kalas & Jewitt (1995) noted that the two extensions of the composite disk are inclined to each other by 1.3° and that the linear fits for

TABLE 3
CHROMATIC VARIATIONS

Parameter	606 χ^2 minimum	606 MCMC best-fit	435 χ^2 minimum	435 MCMC best-fit	814 χ^2 minimum	814 MCMC best-fit
χ^2	6.664×10^5	$6.664^{+0.006}_{-0.002} \times 10^5$	5.395×10^5	$5.396^{+0.003}_{-0.004} \times 10^5$	6.373×10^5	$6.373^{+0.008}_{-0.001} \times 10^5$
Reduced χ^2	1.1833	$1.1834^{+0.0010}_{-0.0003}$	0.9581	$0.9582^{+0.0005}_{-0.0008}$	1.1317	$1.1318^{+0.0014}_{-0.0002}$
i_1 ($^\circ$)	0.11	$0.12^{+0.06}_{-0.06}$	-0.05	$-0.05^{+0.04}_{-0.04}$	0.008	$0.008^{+0.035}_{-0.032}$
p_1	0.84	$0.85^{+0.01}_{-0.02}$	0.82	$0.82^{+0.02}_{-0.02}$	0.83	$0.83^{+0.02}_{-0.01}$
ϕ_1 ($^\circ$)	0.506	$0.505^{+0.028}_{-0.006}$	0.46	$0.46^{+0.02}_{-0.02}$	0.42	$0.42^{+0.02}_{-0.01}$
α_1	2.73	$2.73^{+0.12}_{-0.08}$	3.1	$3.1^{+0.1}_{-0.2}$	2.8	$2.8^{+0.1}_{-0.1}$
β_1	-2.89	$-2.89^{+0.07}_{-0.04}$	-3.09	$-3.10^{+0.06}_{-0.08}$	-2.97	$-2.97^{+0.07}_{-0.06}$
γ_1	1.05	$1.05^{+0.04}_{-0.03}$	0.94	$0.94^{+0.04}_{-0.05}$	0.95	$0.95^{+0.06}_{-0.03}$
i_2 ($^\circ$)	6.312	$6.312^{+0.033}_{-0.010}$	4.88	$4.86^{+0.01}_{-0.05}$	6.766	$6.760^{+0.075}_{-0.006}$
p_2	0.464	$0.465^{+0.024}_{-0.008}$	0.49	$0.49^{+0.02}_{-0.02}$	0.450	$0.453^{+0.029}_{-0.008}$
ϕ_2 ($^\circ$)	3.871	$3.871^{+0.003}_{-0.003}$	4.118	$4.118^{+0.004}_{-0.004}$	2.968	$2.968^{+0.006}_{-0.006}$
α_2	1.6	$1.6^{+0.3}_{-0.3}$	13	12^{+4}_{-10}	0.9	$0.9^{+0.3}_{-0.3}$
β_2	-1.16	$-1.16^{+0.02}_{-0.03}$	-0.95	$-0.95^{+0.02}_{-0.03}$	-1.21	$-1.20^{+0.03}_{-0.04}$
γ_2	0.84	$0.84^{+0.05}_{-0.02}$	0.98	$0.97^{+0.05}_{-0.07}$	0.83	$0.84^{+0.03}_{-0.04}$
R_{m1} (AU)	112.7	$112.7^{+0.9}_{-1.2}$	112.1	$112.2^{+0.8}_{-1.5}$	113	113^{+1}_{-1}
R_{m2} (AU)	71	71^{+3}_{-3}	54	54^{+3}_{-2}	74	72^{+8}_{-2}
η_1	5.60×10^{-10}	$5.62^{+0.06}_{-0.06} \times 10^{-10}$	5.33×10^{-10}	$5.34^{+0.24}_{-0.09} \times 10^{-10}$	5.83×10^{-10}	$5.85^{+0.08}_{-0.08} \times 10^{-10}$
η_2	2.0×10^{-10}	$2.0^{+0.2}_{-0.1} \times 10^{-10}$	2.1×10^{-10}	$2.2^{+0.2}_{-0.1} \times 10^{-10}$	2.2×10^{-10}	$2.2^{+0.4}_{-0.3} \times 10^{-10}$
A_{01}	0.0489	$0.0492^{+0.0010}_{-0.0012}$	0.049	$0.049^{+0.002}_{-0.002}$	0.046	$0.046^{+0.001}_{-0.001}$
A_{02}	0.051	$0.052^{+0.012}_{-0.005}$	0.054	$0.056^{+0.009}_{-0.010}$	0.036	$0.037^{+0.010}_{-0.003}$

^a As discussed in §2.4.1 the formal errors have been scaled up by a factor of 10, as the uncertainties presented here are believed to be a more accurate determination of the uncertainty in the fit parameters.

TABLE 4
RELATIVE POSITION ANGLES OF THE PRIMARY, SECONDARY
AND THE COMPOSITE DISK

Extension	F435W	F606W	F814W	Average
Primary NE	0.454	0.478	0.413	0.448 ± 0.033
Primary SW	0.460	0.519	0.430	0.470 ± 0.045
Secondary NE	2.567	2.400	2.929	2.632 ± 0.270
Secondary SW	5.342	5.006	4.003	4.784 ± 0.697
Composite NE	0.643	0.565	0.446	0.551 ± 0.099
Composite SW	0.880	0.782	0.707	0.789 ± 0.086

* All values given in degrees and measured counterclockwise from the major projection axis. The fits are obtained from regions used by Golimowski et al. (2006), namely 80-250 AU for the primary and the composite and 80-130 AU for the secondary.

their midplanes do not intersect at β Pic. This asymmetry was named the wing-tilt asymmetry and is attributed to forward scattering from an inclined disk. It is also present in the Golimowski et al. (2006) images, where the inclination of the primary's components is determined to be 0.9° while that of the secondary's is 0.3° . Also, they note that the primary's linear fits intersect at a point almost coincident with β Pic, which they note as a point of disagreement.

Our fits also reproduce this asymmetry (Table 4). The inclination of the primary's extensions ($\sim 0.1^\circ$), however, is much less than that reported by Golimowski et al. (2006), but is consistent with a disk that is edge-on as our model suggests the primary is. The linear fits of the disk's spine are nearly centered on the star (Figure 7) in agreement with Golimowski et al. (2006). The secondary's extensions on the other hand are inclined at $\sim 2.2^\circ$ due to the inclination of this component to the

line of sight ($i_2 = 6.0 \pm 1.0^\circ$). Their fits also intersect close to β Pictoris. The composite is inclined at $\sim 0.3^\circ$, which is less than the inclination of the composite determined by Kalas & Jewitt (1995). However, the two extensions do not meet at the star, with the NE component intersecting the SW component ~ 50 AU away from β Pic along the SW extension. The fact that the model primary and composite do not agree on the point of intersection is offered as a possible reason for the disagreement between Kalas & Jewitt (1995) and Golimowski et al. (2006). Our model suggests that it is the combination of the primary and the secondary that together creates the mentioned offset.

Width Asymmetry: Golimowski et al. (2006) find that the FWHM and the full-width-0.1-max (FW0.1M) of the two extensions of the composite diverge beyond ~ 190 AU and ~ 150 AU respectively. This asymmetry, known as the width asymmetry, is present in our fits as well, although to a smaller degree. Figure 11 shows the width of the NE and SW components of the composite. It is evident that the SW component's FWHM is slightly wider than that of the NE component, although not by a significant amount. Out to ~ 200 AU the two extensions FWHM and FW0.1M almost overlap, as is the case with the Golimowski et al. (2006) measurements, but after this point our model disagrees significantly as the Golimowski et al. (2006) measurements show much more asymmetric behaviour. The FW0.1M shows more asymmetric behaviour than the FWHM, as is the case with the Golimowski et al. (2006) measurements.

Surface Brightness Asymmetry: Differences in the brightness profile along the spine of the NE and the SW extensions of the composite are presented in Figure 12. Although our model shows asymmetry between the pro-

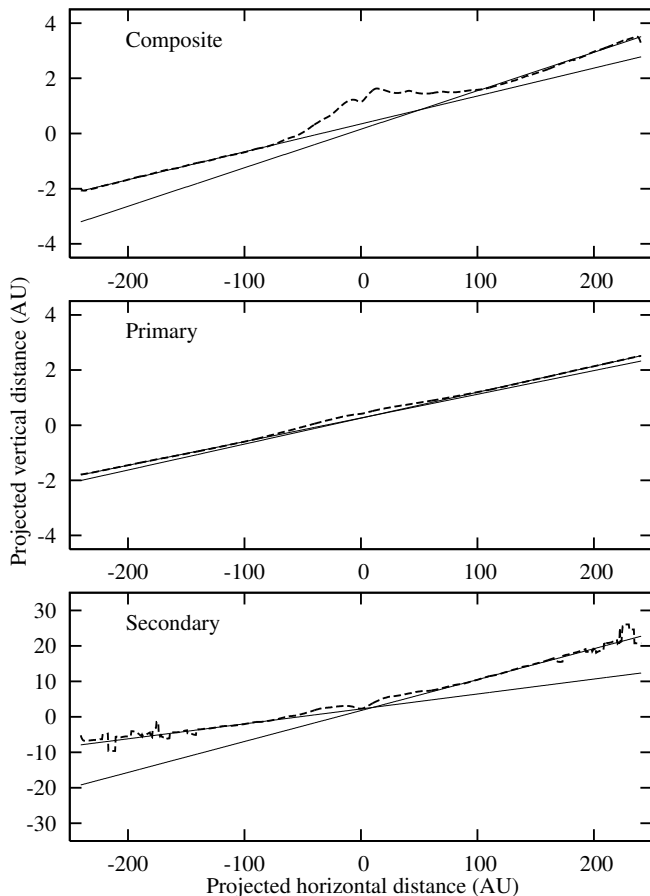


FIG. 7.— Spines of the composite, primary and secondary disks (dashed lines) determined by fitting gaussians to the vertical profiles at each step along the semimajor axis. The solid lines represent linear fits to each extension of the disk, which we use to determine the inclination values of each component presented in Table 4. The solid lines are elongated into the opposite extension for visibility.

files of the two extensions, it does this to a much lesser degree than the scattered light.

Radial Extent Asymmetry: Although the radial extent asymmetry could result from the surface brightness asymmetry, Kalas & Jewitt (1995) classify it separately to draw attention to the possibility that the SW extension is physically truncated. This asymmetry, however, becomes evident only beyond ~ 400 AU and thus we are unable to comment on its true nature.

4.3. Dust Disk Dynamics and Origin

In previous sections we presented a case for two axisymmetric disks around β Pic. Here we explore the dynamical consequences of such a pair. A quantity of interest is the mean collision time since it reveals the timescales on which the disks are stable, and also on which they have to be replenished. It can be estimated most easily by using our knowledge about the total stellar flux absorbed by dust and re-radiated in infrared, quantified as the ratio of IR excess to star’s luminosity $f_d = L_d/L_* = 2.4 \times 10^{-3}$ (e.g., Artymowicz 1997). If the absorbing area of grains in a small volume of space is denoted as dA_{grain} , then we can normalize our model-

derived vertical optical depth functions $\tau(r)$ using

$$f_d = \int \frac{\frac{L_*}{4\pi r^2} dA_{\text{grain}}}{L_*} = \int \frac{k\tau(r)2\pi r dr}{4\pi r^2} = \frac{1}{2} \int \frac{k\tau(r)dr}{r}. \quad (10)$$

We obtain the normalization constants k for each disk separately, utilizing the ratio of scattered light in the two disks ($L_{d,\text{sec}}/L_{d,\text{prim}} = 1.23$ from our models) and the fact that f_d is due to the sum of the two disks. The normalization factors k_{primary} and $k_{\text{secondary}}$ for the two disks can then be used separately or as a sum, to discuss the behavior of individual disks or their sum. In fact, as there is a large degree of overlap between the disks, the second perspective is more realistic. We estimate the mean collision time (t_c) following (Artymowicz 1997) as

$$t_c = \frac{P(r)}{12\tau(r)} \quad (11)$$

where $P(r)$ is the Keplerian period of a particle at radius r . The factor 12 in denominator accounts for the motion parallel to the disk and the fact that the cross section for collision between two particles of comparable radius is up to 4 times bigger than the geometrical cross section of either one.

In Figure 13 we present the conservative number of orbits (assuming the disks completely overlap) a particle can make before colliding with another particle as a function of distance, and the corresponding mean collisional time scale.

Since the lifetime of dust in the densest parts of the disks is of order 10^4 to 10^5 yr, and hence much less than that of the host star (which is at least 20 Myr old), our calculations suggest both disks have planetesimals embedded in them capable of replenishing the lost dust. In other words, none of the disks represent primordial protoplanetary disks. Rather, they are replenished disk(s) owing their existence to a substantial mass reservoir, equivalent to a planetary system comparable to our own (Artymowicz 1997).

Geometrically, the two disks largely overlap and some regions contain comparable areas of dust. Physically, this implies mutual influence through collisions. Since the optical depths of the disks are much smaller than unity, particles in the two disks can coexist for as many as 20-10000 of their orbits. Similarly, the smaller optical depth of the larger parent bodies assures their stability for a large number of orbits in the overlapping disks. The inferred presence of two disks suggests that β Pic has two distinct populations of parent bodies (planetesimals) in a slightly inclined configuration.

As mentioned in the introduction, the characterization of the distribution of dust around a star is but a first step toward the derivation of the mass and spatial distribution of larger, unseen bodies (planetesimals), as well as any perturbing planets; other effects such as dust-dust and any gas-dust interaction in the presence of stellar radiation pressure can spatially displace the fine dust from the place of their origin (the parent bodies). However, some implications of our models are already apparent.

The fact that the secondary disk contains a significant fraction of the dust in the system means that its origin cannot be ascribed to a minor event such as a single comet or an asteroid disruption. In order to produce comparable amounts of dust in the two disks, they must

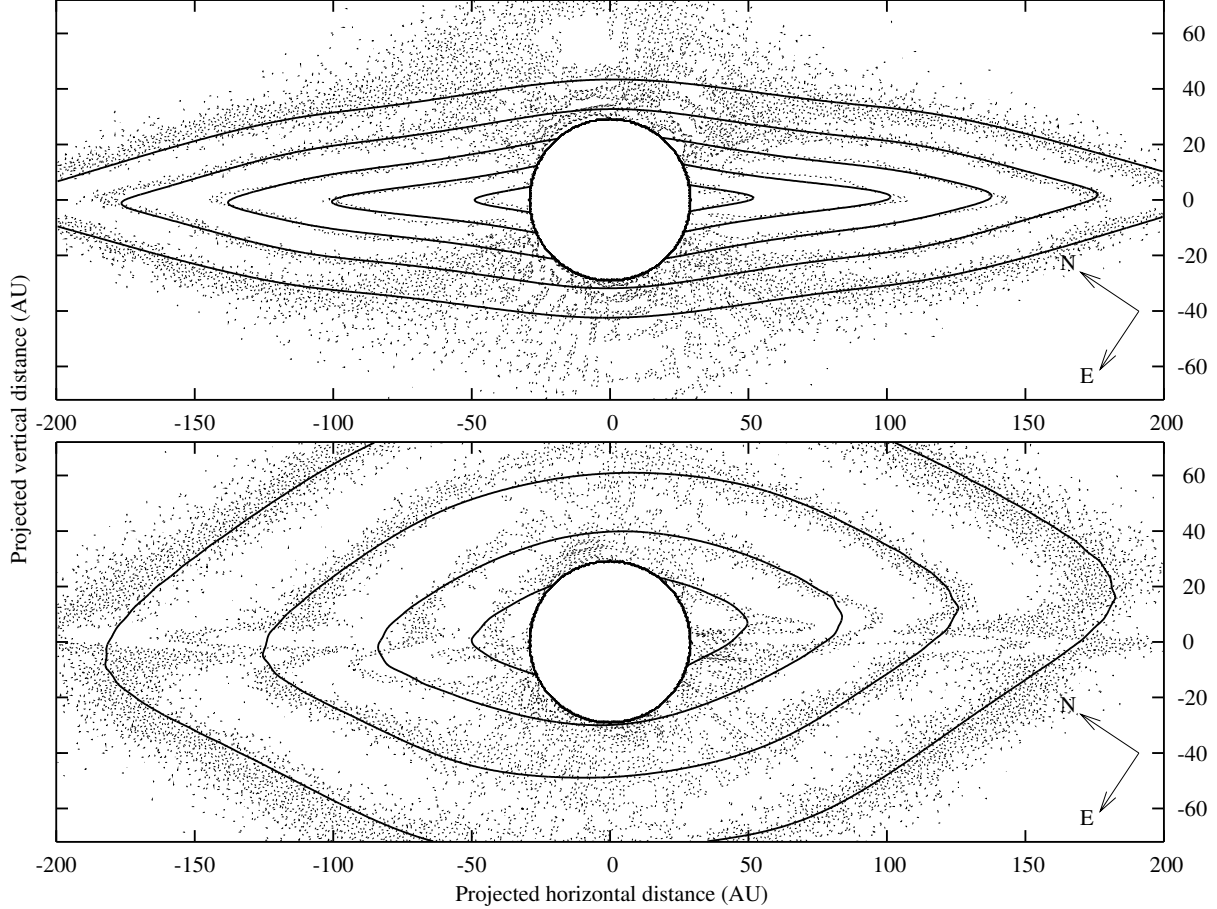


FIG. 8.— Isophotal maps of light scattered by the primary (top panel) and the secondary (bottom panel). Solid lines are the isophotes of the models while the grey regions underneath represent residuals following the removal of the secondary (top panel) and the primary (bottom panel) fits. The brightness isophotes are quoted relative to βPic , with the isophotes decreasing from $1 \times 10^{-7} L_{disk}/L_{\beta Pic}$ in the case of the primary and from $3.162 \times 10^{-7} L_{disk}/L_{\beta Pic}$ in the case of the secondary. In both panels the spacing of the isophotes is equal to 0.5 in log base 10 units. The empty circular region in the centre of the image represents the coronagraphic spot covering this area. The most important features of these fits include 1. extended secondary which contributes more light than the primary at larger radii and scale heights, 2. more widely spaced isophotes to the NW than SE of the secondary demonstrating significant inclination and 3. nearly edge-on ($i_1 = 0.0$) primary.

either have a common massive reservoir of parent bodies to draw from or two such reservoirs. Only comprehensive modelling based on high-quality data can clarify the situation since the intermediate-size bodies such as km-sized planetesimals are not amenable to direct detection, possessing neither enough gravitational influence nor large emitting surface area.

The vertical profile we obtained is $\rho \sim e^{-(|z|/W(r))^p}$, with $p = p_1 \approx 0.83$ for the first disk and $p = p_2 \approx 0.46$ for the second. Thus, neither disk has a vertical scattering profile resembling a disk with a vertical isothermal structure ($p = 2$, or Gaussian vertical profile). The two disks, especially the second one, have super-exponential drop-offs ($p < 1$) and therefore extended, slowly diminishing outer wings of the vertical structure, with a sharply pointed inner core. This could be a sign that the inner core of the profile, referred to as the spine in this paper, is composed of bigger particles than the profile wings; the mass separation might arise due to energy equipartition leading to a smaller velocity dispersion of the larger particles in the disk. In addition, a more fanned out secondary debris disk could be a consequence of more en-

ergetic collisions of its particles, perhaps indicating their generally smaller sizes and larger radiative modification of orbits.

An expectation based on the the previous dynamical studies of the mobility of dust is that the bulk of a disk or a ring of planetesimals will be located interior to the dust disk. This follows from the outward action of radiation pressure exerted on dust. If so, we expect that most planetesimals will be found at a distances less than 50-100 AU from βPic . At such distances, planetesimals can be perturbed by the proposed 8-Jupiter mass planet at a mean distance of 8 AU and/or other unknown planetary masses. Perturbing planets, in any case, must be located inside the cross-over radius ($R_{m2} < 70$ AU) of the secondary disk.

We stress here the possibility, indeed necessity, of multiple planets in βPic . A planetary system born with only one major planet would tend to stay in alignment with the orbital plane of this planet. In such a case, we would not expect to see two separate inclined disks. What could be the origin of the inclined orbit of an additional planet? Although specific inclination pumping

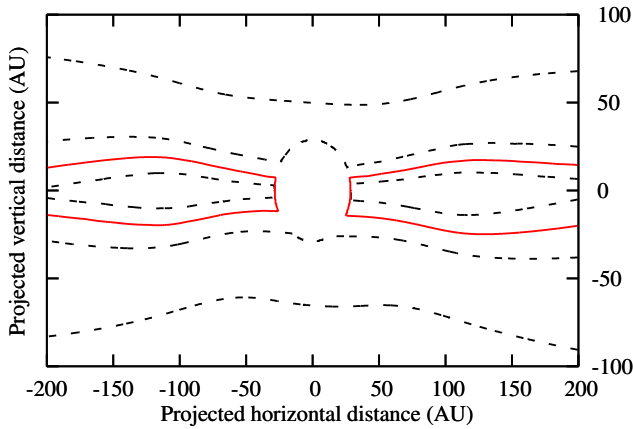


FIG. 9.— Ratio of the projected light scattered by the two disks, demonstrating that along the midplane light from the primary dominates, while although at larger scale heights the secondary becomes dominant. The solid line outlines the region inside of which the primary's light dominates. The dashed line inside the solid line denotes the isophote where the primary's contribution is twice that of the secondary. The dashed lines outside the region represent isophotes where the secondary's contribution is twice and ten times greater than the primary's, the latter being the isophote at larger scale height.

mechanisms have yet to be quantified, this may provide a natural explanation to the inclination of planetesimals and dust from the second disk. This could result in the planet(s) themselves being on slightly inclined, mutually interacting, orbits thus contributing to the survival of the two inclined disks.

Future observations of the inner regions of the disks would be instrumental in establishing the presence of the putative multiple planets via their direct influence on dust. At some point, the assumption made in the present work that the disks are intrinsically axisymmetric, if misaligned, must prove inadequate. With sufficient observational data, our assumption of the same optical properties of dust in two disks may also be relaxed. Furthermore, a future comprehensive model of the system must include not only collisional dust dynamics and radiation, but also the collisional cascades spanning a large size range from planetesimals to micron-sized dust, and possibly more unusual ingredients, such as dust avalanches involving sub-micron dust.

5. CONCLUSIONS

Below we summarize the results of this work.

1. We modeled the dust distribution around β Pic with two axisymmetric disks. A MCMC χ^2 optimization method was used to reject disk models and phase functions that produce brightness contours different from those obtained with the ACS by Golimowski et al. (2006). We find a small quantitative improvement in the reduced χ^2 (from 1.8 to 1.2) of the two-disk fit compared to the single disk fit; the two-disk fit, however, displays significant improvements in its ability to trace the isophotes of the composite. The parameters listed in Table 3 along with a zodiacal light scattering profile produce the best fits.

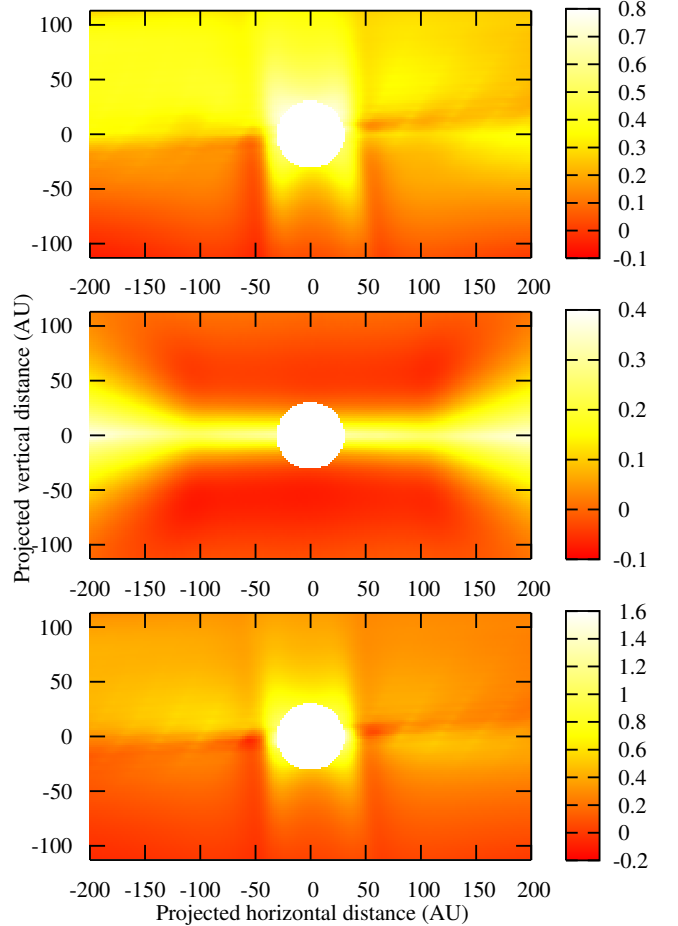


FIG. 10.— Ratios of light scattered by the composite (top panel), primary (middle panel) and secondary (bottom panel) disks at 435 and 814 nm. The red region shows the highest relative contribution in the 814 nm relative to the 435 nm filter. The ratios are presented in log base 2 units. From the presented ratios, we conclude that the color asymmetries along the projected major and minor axis are due to an inclined secondary disk.

2. Dust distribution in the primary disk is largely in agreement with conclusions from previous modeling attempts (eg. Artymowicz et al. 1989; Kalas & Jewitt 1995). The modelled primary disk has a near-zero inclination to the line of sight. Therefore, the up-down asymmetry of the whole disk is not caused by the primary. The surface density of dust increases as $\alpha_1 = 2.73$ -th power of radius inside a cross-over radius $R_{m1} = 113$ AU and decreases at large radii with power-law index $\beta_1 = -2.89$ (values from F606W filter fit). This behavior is similar to the one known from previous modelling, although the inner power law is steeper. The vertical disk profile has a nearly exponential fall-off with vertical distance ($p_1 = 0.85$). The scattered light contours from the primary disk are shown in Figure 8.
3. The secondary disk model presents the first quantitative measurement of the dust distribution in this component. It suggests that the secondary disk, with a scale-height much larger than the primary, is not radially limited to the inner system,

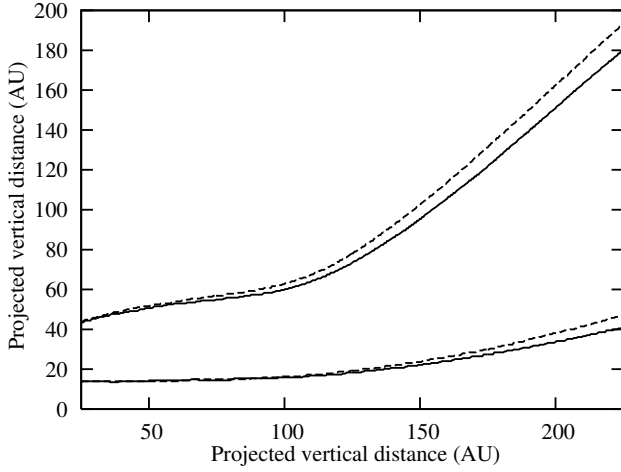


FIG. 11.— FWHM (lower pair) and the FW0.1M (upper pair) of our model for the NE (solid lines) and SW (dashed lines). The wider SW extension demonstrates that our model reproduces the width asymmetry first noted by Kalas & Jewitt (1995).

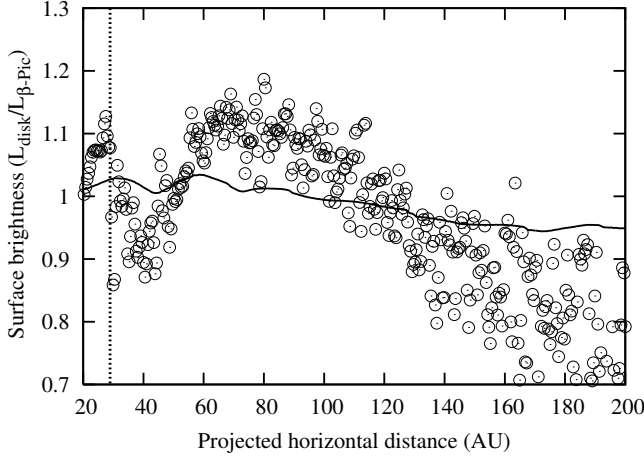


FIG. 12.— The solid line represents the ratio of the brightness profile along the midplane of the SW extension divided by the corresponding brightness profile of the NE extension of our best-fit 18-parameter model to the 606 ACS data. The same ratio is shown for the 606 ACS scattered light data (open circles). All open circles represent a 3 pixel average taken along the minor projection axis (the y-axis of Figure 4). The model profile along the spines of the two extensions shows asymmetric behaviour, but the degree of this asymmetry is far less than that observed from the scattered light. The dotted vertical line at 30 AU represents the extent of the coronagraphic mask.

but that it extends at least to 250 AU, and possibly further out. The secondary disk model has a larger scattering area than the primary at large distances and scale heights, which follows from the less-steep power laws of the radial surface density ($\beta_2 = -1.16 > -2.89$ at distances much larger than the cross-over $R_{m2} = 71$ AU) and the vertical density ($p_2 = 0.46 < 0.85$). The ratio of the total scattered light (secondary vs. primary) is roughly equal to 1.23. Since we do not model the dust size distribution, we cannot compute the dust mass ra-

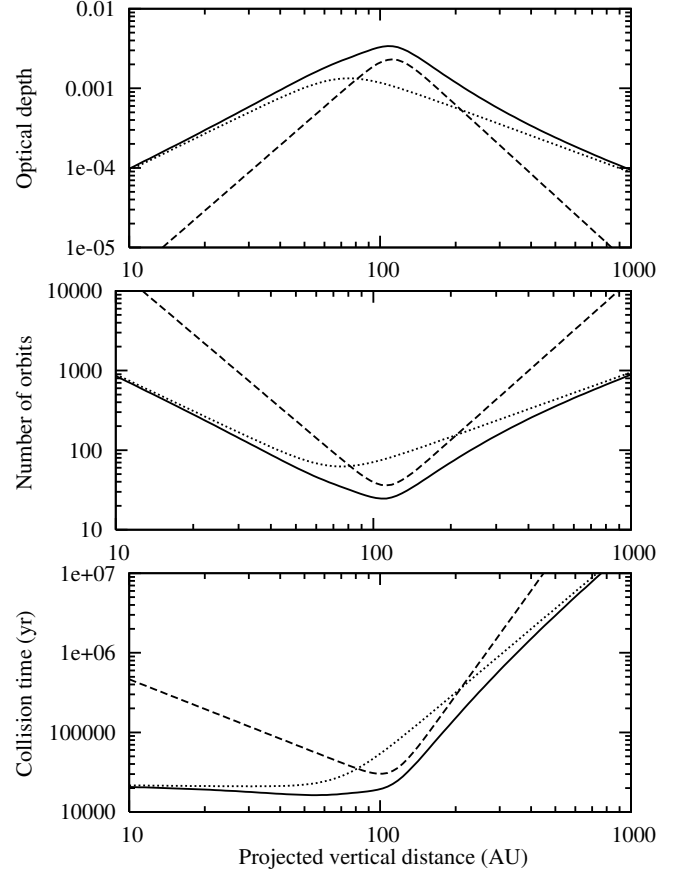


FIG. 13.— The optical depths and mean collisional times for the two disks (assuming they are completely physically separated; dashed lines for the primary, dotted lines for the secondary) as well as the conservative estimate (assuming that the two disks overlap completely; solid lines). Due to the uncertainty in our model of the inner radial index (α_2), the collisional times for the inner regions of the secondary disk ($< \sim 60$ AU) are less reliable.

tio of the two components. It is however plausible that the secondary disk is as substantial as the primary. The larger extent and especially the larger scattering area of the secondary disk are surprising and raise many questions about the origin and dynamics of such a pair of disks. The scattered light contours from the secondary are presented in Figure 8.

4. Several of the asymmetries reported by Kalas & Jewitt (1995) can be reproduced with our model. The butterfly asymmetry is explained by the presence of an extended secondary disk inclined by $3.2 \pm 1.3^\circ$ to the primary. We find that it is the secondary that is inclined to the line of sight and therefore responsible for the wing-tilt asymmetry observed in the composite image. The width asymmetry is present, but not to the degree seen in Golimowski et al. (2006) data beyond ~ 200 AU. The brightness asymmetry is also noticeable both across the major and the minor projection axis. The difference across both the axes is attributed to a secondary disk inclined

by $i_2 = 6.0 \pm 1.0^\circ$ to the line of sight and to the primary which is nearly aligned with the major projection axis.

5. In order to explore dust dynamics in a pair of disks described by the parameters obtained here, we analytically constrain the dust replenishment times to be 10^4 yr at ~ 100 AU. Such small timescales show that both disks have to be continuously replenished and contain planetesimals that, as eventual mass reservoir, are capable of producing the second-generation dust. The nature and origin of the inclined disks and planetesimals in β Pic should be explored via dynamical models of spatially resolved collisional cascades subject to radiation pressure effects.
6. The most plausible conjectural outcome of this modeling is that the non-coplanarity of two dust

disks is due to the secondary dust disk being supplied by its own disk of planetesimals, which has been inclined by the gravity of a planet residing inside the cross-over radii ($r < 70$ AU). This, in turn, would be most naturally explained if there were more than one massive planets in the system, in orbits inclined by several degrees. Future modeling of unseen planetesimals of β Pictoris, constrained by the new knowledge of dust distribution, may clarify this conjecture.

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